

AperTO - Archivio Istituzionale Open Access dell'Università di Torino

Cosmic ray oriented performance studies for the JEM-EUSO first level trigger

This is a pre print version of the following article:

Original Citation:

Availability:

This version is available <http://hdl.handle.net/2318/1653095> since 2017-11-25T11:40:01Z

Published version:

DOI:10.1016/j.nima.2017.05.043

Terms of use:

Open Access

Anyone can freely access the full text of works made available as "Open Access". Works made available under a Creative Commons license can be used according to the terms and conditions of said license. Use of all other works requires consent of the right holder (author or publisher) if not exempted from copyright protection by the applicable law.

(Article begins on next page)

Cosmic Ray Oriented Performance Studies for the JEM-EUSO First Level Trigger

G. Abdellaoui^{ah}, S. Abe^{fu}, A. Acheli^{aa}, J.H. Adams Jr.^{pd}, S. Ahmad^{cb},
A. Ahriche^{ae}, J.-N. Albert^{ca}, D. Allard^{cc}, G. Alonso^{md}, L. Anchordoqui^{pf},
V. Andreev^{pe}, A. Anzalone^{eh,en}, W. Aouimeur^{aa}, Y. Arai^{fw}, N. Arsene^{ja},
K. Asano^{fg}, R. Attallah^{ac}, H. Attoui^{aa}, M. Ave Pernas^{mc}, S. Bacholle^{cc},
M. Bakiri^{aa}, P. Baragatti^{eo}, P. Barrillon^{ca}, S. Bartocci^{eo}, T. Batsch^{ic},
J. Bayer^{dd}, R. Bechini^{el}, T. Belenguer^{mb}, R. Bellotti^{ea,eb}, A. Belov^{kc},
K. Belov^{pe}, B. Benadda^{ah}, K. Benmessai^{ag}, A.A. Berlind^{ph},
M. Bertaina^{ek,el,1}, P.L. Biermann^{db}, S. Biktemerova^{ka}, F. Bisconti^{db},
N. Blanc^{oa}, J. Błęcki^{id}, S. Blin-Bondil^{cb}, P. Bobik^{la}, M. Bogomilov^{ba},
M. Bonamente^{pd}, R. Boudaoud^{aa}, E. Bozzo^{ob}, M.S. Briggs^{pd}, A. Bruno^{eb},
K.S. Caballero^{he}, F. Cafagna^{ea}, D. Campana^{ef}, J.-N. Capdevielle^{cc},
F. Capel^{na}, A. Caramete^{ja}, L. Caramete^{ja}, P. Carlson^{na}, R. Caruso^{ec,en,1},
M. Casolino^{fx,ei}, C. Cassardo^{ek,el}, A. Castellina^{ek,em}, G. Castellini^{ed},
C. Catalano^{cd}, O. Catalano^{eh,en,1}, A. Cellino^{ek,em}, M. Chikawa^{fd},
G. Chiritoi^{ja}, M.J. Christl^{pg}, V. Connaughton^{pd}, L. Conti^{eo}, G. Contino^{ec,en},
G. Cordero^{ha}, G. Cotto^{ek,el}, H.J. Crawford^{pa}, R. Cremonini^{el}, S. Csorna^{ph},
S. Dagoret-Campagne^{ca}, C. De Donato^{ei}, C. de la Taille^{cb}, C. De Santis^{ei},
L. del Peral^{mc}, M. Di Martino^{em}, T. Djemil^{ac}, S.A. Djenas^{ah}, F. Dulucq^{cb},
M. Dupieux^{cd}, I. Dutan^{ja}, A. Ebersoldt^{db}, T. Ebisuzaki^{fx}, R. Engel^{db},
J. Eser^{pc}, K. Fang^{pb}, F. Fenu^{ek,el}, S. Fernández-González^{ma},
J. Fernández-Soriano^{mc}, S. Ferrarese^{ek,el}, D. Finco^{eo}, M. Flamini^{eo},
C. Fornaro^{eo}, R. Forza^{ek,el}, M. Fouka^{ab}, A. Franceschi^{ee}, S. Franchini^{md},
C. Fuglesang^{na}, J. Fujimoto^{fw}, M. Fukushima^{fg}, P. Galeotti^{ek,el},
E. García-Ortega^{ma}, G. Garipov^{kc}, E. Gascón^{ma}, J. Geary^{pd}, G. Gelmini^{pe},
J. Genci^{lb}, G. Giraudo^{ek}, M. Gonchar^{ka}, C. González Alvarado^{mb},
P. Gorodetzky^{cc}, N. Guardone^{ec,en}, F. Guarino^{ef,eg}, R. Guehaz^{aa},
A. Guzmán^{dd}, Y. Hachisu^{fx}, M. Haiduc^{ja}, B. Harlov^{kb}, A. Haungs^{db},
J. Hernández Carretero^{mc}, W. Hidber^{ha}, K. Higashide^{fr,fx}, D. Ikeda^{fg},
H. Ikeda^{fp}, N. Inoue^{fr}, S. Inoue^{fx}, A. Insolia^{ec,en}, F. Isgrò^{ef,ep}, Y. Itow^{fn},
T. Jammer^{dc}, E. Joven^{me}, E.G. Judd^{pa}, A. Jung^{cc}, J. Jochum^{dc}, F. Kajino^{fi},
T. Kajino^{fl}, S. Kalli^{af}, I. Kaneko^{fx}, D. Kang^{db}, F. Kanouni^{ag},
Y. Karadzhov^{ba}, J. Karczmarczyk^{ic}, M. Karus^{db}, K. Katahira^{fx},
K. Kawai^{fx}, Y. Kawasaki^{fx}, A. Kedadra^{aa}, H. Kholes^{aa}, B.A. Khrenov^{kc},
Jeong-Sook Kim^{ga}, Soon-Wook Kim^{ga}, Sug-Whan Kim^{gd}, M. Kleifges^{db},

P.A. Klimov^{kc}, D. Kolev^{ba}, I. Kreykenbohm^{da}, K. Kudela^{la}, Y. Kurihara^{fw},
 A. Kusenko^{fv,pe}, E. Kuznetsov^{pd}, M. Lacombe^{cd}, C. Lachaud^{cc},
 H. Lahmar^{aa}, F. Lakhdari^{ag}, O. Larsson^{fx,na}, J. Lee^{gc}, J. Licandro^{me},
 H. Lim^{gc}, L. López Campano^{ma}, M.C. Maccarone^{eh,en}, S. Mackovjak^{ob},
 M. Mahdi^{aa}, M. Manfrin^{ek,el}, D. Maravilla^{ha}, L. Marcelli^{ej}, J.L. Marcos^{ma},
 A. Marini^{ee}, K. Martens^{fv}, Y. Martín^{me}, O. Martinez^{hc}, G. Masciantonio^{ei},
 K. Mase^{fa}, R. Matev^{ba}, J.N. Matthews^{pi}, N. Mebarki^{ad},
 G. Medina-Tanco^{ha}, L. Mehrad^{ah}, M.A. Mendoza^{hd}, A. Merino^{ma},
 T. Mernik^{dd}, J. Meseguer^{md}, S. Messaoud^{aa}, O. Micu^{ja}, M. Mignone^{ek,el},
 J. Mimouni^{ad}, H. Miyamoto^{ek,el,1}, Y. Miyazaki^{fc}, Y. Mizumoto^{fl},
 G. Modestino^{ee}, A. Monaco^{ea,eb}, D. Monnier-Ragaigne^{ca}, J.A. Morales de
 los Ríos^{mc}, C. Moretto^{ca}, V.S. Morozenko^{kc}, B. Mot^{cd}, T. Murakami^{ff},
 B. Nadji^{aa}, M. Nagano^{fc}, M. Nagata^{fh}, S. Nagataki^{fx}, T. Nakamura^{fj},
 T. Napolitano^{ee}, D. Naumov^{ka}, R. Nava^{ha}, A. Neronov^{ob}, K. Nomoto^{fv},
 T. Nonaka^{fg}, T. Ogawa^{fx}, S. Ogio^{fo}, H. Ohmori^{fx}, A.V. Olinto^{pb},
 P. Orleañski^{id}, G. Osteria^{ef}, W. Painter^{db}, M.I. Panasyuk^{kc}, B. Panico^{ef},
 E. Parizot^{cc}, I.H. Park^{gc}, H.W. Park^{gc}, B. Pastircak^{la}, T. Patzak^{cc},
 T. Paul^{pf}, C. Pennypacker^{pa}, I. Pérez-Grande^{md}, F. Perfetto^{ef,eg}, T. Peter^{oc},
 P. Picozza^{ei,ej,fx}, T. Pierog^{db}, S. Pindado^{md}, L.W. Piotrowski^{pc},
 S. Piraino^{dd,eh}, L. Placidi^{eo}, Z. Plebaniak^{ic}, S. Pliego^{ha}, A. Pollini^{oa},
 E.M. Popescu^{ja}, P. Prat^{cc}, G. Prévôt^{cc}, H. Prieto^{mc}, M. Putis^{la},
 J. Rabanal^{ca}, A.A. Radu^{ja}, M. Rahmani^{ag}, P. Reardon^{pd}, M. Reyes^{me},
 M. Rezazadeh^{pb}, M. Ricci^{ee}, M.D. Rodríguez Frías^{mc}, F. Ronga^{ee},
 M. Roth^{db}, H. Rothkaehl^{id}, G. Roudil^{cd}, I. Rusinov^{ba}, M. Rybczyński^{ia},
 M.D. Sabau^{mb}, G. Sáez Cano^{mc}, H. Sagawa^{fg}, Z. Sahnoun^{ab}, A. Saito^{fj},
 N. Sakaki^{fo}, M. Sakata^{fi}, H. Salazar^{hc}, J.C. Sanchez^{ha}, J.L. Sánchez^{ma},
 A. Santangelo^{dd}, L. Santiago Crúz^{ha}, A. Sanz-Andrés^{md},
 M. Sanz Palomino^{mb}, O. Saprykin^{kb}, F. Sarazin^{pc}, H. Sato^{fi}, M. Sato^{fs},
 T. Schanz^{dd}, H. Schieler^{db}, V. Scotti^{ef,eg}, A. Segreto^{eh,en}, S. Selmane^{cc},
 D. Semikoz^{cc}, M. Serra^{me}, S. Sharakin^{kc}, T. Shibata^{fq}, H.M. Shimizu^{fm},
 K. Shinozaki^{dd}, T. Shirahama^{fr}, G. Siemienieć-Oziębło^{ib}, J. Sledd^{pg},
 K. Słomińska^{id}, A. Sobey^{pg}, I. Stan^{ja}, T. Sugiyama^{fm}, D. Supanitsky^{ha},
 M. Suzuki^{fp}, B. Szabelska^{ic}, J. Szabelski^{ic}, H. Tahia^{aa}, F. Tajima^{fe},
 N. Tajima^{fx}, T. Tajima^{fx}, Y. Takahashi^{fs}, H. Takami^{fw}, M. Takeda^{fg},
 Y. Takizawa^{fx}, M.C. Talai^{ac}, C. Tenzer^{dd}, O. Tibolla^{hf}, L. Tkachev^{ka},
 H. Tokuno^{ft}, T. Tomida^{fk}, N. Tone^{fx}, S. Toscano^{ob}, M. Traïche^{aa},
 R. Tsenov^{ba}, Y. Tsunesada^{fo}, K. Tsuno^{fx}, T. Tymieniecka^{ic}, Y. Uchihori^{fb},
 M. Unger^{db}, O. Vaduvescu^{me}, J.F. Valdés-Galicia^{ha}, P. Vallania^{ek,em},

G. Vankova^{ba}, C. Vigorito^{ek,el}, L. Villaseñor^{hb}, B. Vlcek^{mc}, P. von Ballmoos^{cd}, M. Vrabel^{lb}, S. Wada^{fx}, J. Watanabe^{fl}, S. Watanabe^{fs}, J. Watts Jr.^{pd}, M. Weber^{db}, R. Weigand Muñoz^{ma}, A. Weindl^{db}, T.J. Weiler^{ph}, T. Wibig^{ic}, L. Wiencke^{pc}, M. Wille^{da}, J. Wilms^{da}, Z. Włodarczyk^{ia}, T. Yamamoto^{fi}, Y. Yamamoto^{fi}, J. Yang^{gb}, H. Yano^{fp}, I.V. Yashin^{kc}, D. Yonetoku^{ff}, S. Yoshida^{fa}, R. Young^{pg}, I.S. Zgura^{ja}, M.Yu. Zotov^{kc}, A. Zuccaro Marchi^{fx}

- ^{aa} Centre for Development of Advanced Technologies (CDTA), Algiers, Algeria
^{ab} Dep. Astronomy, Centre Res. Astronomy, Astrophysics and Geophysics (CRAAG), Algiers, Algeria
^{ac} LPR at Dept. of Physics, Faculty of Sciences, University Badji Mokhtar, Annaba, Algeria
^{ad} Lab. of Math. and Sub-Atomic Phys. (LPMPS), Univ. Constantine I, Constantine, Algeria
^{ae} Laboratory of Theoretical Physics LPT, University of Jijel, Jijel, Algeria
^{af} Department of Physics, Faculty of Sciences, University of M'sila, M'sila, Algeria
^{ag} Research Unit on Optics and Photonics, UROP-CDTA, Sétif, Algeria
^{ah} Telecom Lab., Faculty of Technology, University Abou Bekr Belkaid, Tlemcen, Algeria
^{ba} St. Kliment Ohridski University of Sofia, Bulgaria
^{ca} LAL, Univ Paris-Sud, CNRS/IN2P3, Orsay, France
^{cb} Omega, Ecole Polytechnique, CNRS/IN2P3, Palaiseau, France
^{cc} APC, Univ Paris Diderot, CNRS/IN2P3, CEA/Irfu, Obs de Paris, Sorbonne Paris Cité, France
^{cd} IRAP, Université de Toulouse, CNRS, Toulouse, France
^{da} ECAP, University of Erlangen-Nuremberg, Germany
^{db} Karlsruhe Institute of Technology (KIT), Germany
^{dc} Experimental Physics Institute, Kepler Center, University of Tübingen, Germany
^{dd} Institute for Astronomy and Astrophysics, Kepler Center, University of Tübingen, Germany
^{ea} Istituto Nazionale di Fisica Nucleare - Sezione di Bari, Italy
^{eb} Università degli Studi di Bari Aldo Moro and INFN - Sezione di Bari, Italy
^{ec} Dipartimento di Fisica e Astronomia - Università di Catania, Italy
^{ed} Consiglio Nazionale delle Ricerche (CNR) - Istituto di Fisica Applicata Nello Carrara, Firenze, Italy
^{ee} Istituto Nazionale di Fisica Nucleare - Laboratori Nazionali di Frascati, Italy
^{ef} Istituto Nazionale di Fisica Nucleare - Sezione di Napoli, Italy
^{eg} Università di Napoli Federico II - Dipartimento di Scienze Fisiche, Italy
^{eh} INAF - Istituto di Astrofisica Spaziale e Fisica Cosmica di Palermo, Italy
^{ei} Istituto Nazionale di Fisica Nucleare - Sezione di Roma Tor Vergata, Italy
^{ej} Università di Roma Tor Vergata - Dipartimento di Fisica, Roma, Italy
^{ek} Istituto Nazionale di Fisica Nucleare - Sezione di Torino, Italy
^{el} Dipartimento di Fisica, Università di Torino, Italy
^{em} Osservatorio Astrofisico di Torino, Istituto Nazionale di Astrofisica, Italy

- ^{en} *Istituto Nazionale di Fisica Nucleare - Sezione di Catania, Italy*
- ^{eo} *UTIU, Dipartimento di Ingegneria, Rome, Italy*
- ^{ep} *DIETI, Università degli Studi di Napoli Federico II, Napoli, Italy*
 - ^{fa} *Chiba University, Chiba, Japan*
 - ^{fb} *National Institute of Radiological Sciences, Chiba, Japan*
 - ^{fc} *Fukui University of Technology, Fukui, Japan*
 - ^{fd} *Kinki University, Higashi-Osaka, Japan*
 - ^{fe} *Hiroshima University, Hiroshima, Japan*
 - ^{ff} *Kanazawa University, Kanazawa, Japan*
 - ^{fg} *Institute for Cosmic Ray Research, University of Tokyo, Kashiwa, Japan*
 - ^{fh} *Kobe University, Kobe, Japan*
 - ^{fi} *Konan University, Kobe, Japan*
 - ^{fj} *Kyoto University, Kyoto, Japan*
 - ^{fk} *Shinshu University, Nagano, Japan*
 - ^{fl} *National Astronomical Observatory, Mitaka, Japan*
 - ^{fm} *Nagoya University, Nagoya, Japan*
 - ^{fn} *Institute for Space-Earth Environmental Research, Nagoya University, Nagoya, Japan*
 - ^{fo} *Graduate School of Science, Osaka City University, Japan*
 - ^{fp} *Institute of Space and Astronautical Science/JAXA, Sagamihara, Japan*
 - ^{fq} *Aoyama Gakuin University, Sagamihara, Japan*
 - ^{fr} *Saitama University, Saitama, Japan*
 - ^{fs} *Hokkaido University, Sapporo, Japan*
 - ^{ft} *Interactive Research Center of Science, Tokyo Institute of Technology, Tokyo, Japan*
 - ^{fu} *Nihon University Chiyoda, Tokyo, Japan*
 - ^{fv} *Kavli-IPMU (WPI), The University of Tokyo, Kashiwa, Japan*
 - ^{fw} *High Energy Accelerator Research Organization (KEK), Tsukuba, Japan*
 - ^{fx} *RIKEN, Wako, Japan*
 - ^{ga} *Korea Astronomy and Space Science Institute (KASI), Daejeon, Republic of Korea*
 - ^{gb} *Ewha Womans University, Seoul, Republic of Korea*
 - ^{gc} *Sungkyunkwan University, Seoul, Republic of Korea*
 - ^{gd} *Center for Galaxy Evolution Research, Yonsei University, Seoul, Republic of Korea*
 - ^{ha} *Universidad Nacional Autónoma de México (UNAM), Mexico*
 - ^{hb} *Universidad Michoacana de San Nicolas de Hidalgo (UMSNH), Morelia, Mexico*
 - ^{hc} *Benemérita Universidad Autónoma de Puebla (BUAP), Mexico*
 - ^{hd} *Centro de Desarrollo Aeroespacial - Instituto Politécnico Nacional (CDA-IPN), Mexico*
 - ^{he} *Universidad Autónoma de Chiapas (UNACH), Chiapas, Mexico*
 - ^{hf} *Centro Mesoamericano de Física Teórica (MCTP), Mexico*
 - ^{ia} *Jan Kochanowski University, Institute of Physics, Kielce, Poland*
 - ^{ib} *Jagiellonian University, Astronomical Observatory, Krakow, Poland*
 - ^{ic} *National Centre for Nuclear Research, Lodz, Poland*
 - ^{id} *Space Research Centre of the Polish Academy of Sciences (CBK), Warsaw, Poland*
 - ^{ja} *Institute of Space Science ISS, Magurele, Romania*
 - ^{ka} *Joint Institute for Nuclear Research, Dubna, Russia*
 - ^{kb} *Central Research Institute of Machine Building, TsNIIMash, Korolev, Russia*

- ^{kc} Skobeltsyn Institute of Nuclear Physics, Lomonosov Moscow State University, Russia
^{la} Institute of Experimental Physics, Kosice, Slovakia
^{lb} Technical University Kosice (TUKE), Kosice, Slovakia
^{ma} Universidad de León (ULE), León, Spain
^{mb} Instituto Nacional de Técnica Aeroespacial (INTA), Madrid, Spain
^{mc} Universidad de Alcalá (UAH), Madrid, Spain
^{md} Universidad Politécnica de Madrid (UPM), Madrid, Spain
^{me} Instituto de Astrofísica de Canarias (IAC), Tenerife, Spain
^{na} KTH Royal Institute of Technology, Stockholm, Sweden
^{oa} Swiss Center for Electronics and Microtechnology (CSEM), Neuchâtel, Switzerland
^{ob} ISDC Data Centre for Astrophysics, Versoix, Switzerland
^{oc} Institute for Atmospheric and Climate Science, ETH Zürich, Switzerland
^{pa} Space Science Laboratory, University of California, Berkeley, USA
^{pb} University of Chicago, USA
^{pc} Colorado School of Mines, Golden, USA
^{pd} University of Alabama in Huntsville, Huntsville, USA
^{pe} University of California (UCLA), Los Angeles, USA
^{pf} Lehman College, City University of New York (CUNY), USA
^{pg} NASA - Marshall Space Flight Center, USA
^{ph} Vanderbilt University, Nashville, USA
^{pi} University of Utah, Salt Lake City, USA
- ¹ Corresponding authors: Mario Bertaina (E-mail address: bertaina@to.infn.it), Rossella Caruso (E-mail address: rossella.caruso@ct.infn.it), Osvaldo Catalano (E-mail address: catalano@ifc.inaf.it), Hiroko Miyamoto (E-mail address: miyamoto@to.infn.it)

Abstract

JEM-EUSO is a space mission designed to investigate Ultra-High Energy Cosmic Rays and Neutrinos ($E > 5 \cdot 10^{19}$ eV) from the International Space Station (ISS). Looking down from above its wide angle telescope is able to observe their air showers and collect such data from a very wide area. Highly specific trigger algorithms are needed to drastically reduce the data load in the presence of both atmospheric and human activity related background light, yet retain the rare cosmic ray events recorded in the telescope. We report the performance in offline testing of the first level trigger algorithm on data from JEM-EUSO prototypes and laboratory measurements observing different light sources: data taken during a high altitude balloon flight over Canada, laser pulses observed from the ground traversing the real atmosphere, and model landscapes reproducing realistic aspect ratios and light

conditions as would be seen from the ISS itself. The first level trigger logic successfully kept the trigger rate within the permissible bounds when challenged with artificially produced as well as naturally encountered night sky background fluctuations and while retaining events with general air-shower characteristics.

Key words: JEM-EUSO, trigger system, FPGA, nightglow background

1. Introduction

Ultra-High Energy Cosmic Rays (UHECR) are observed as Extensive Air Showers (EAS) in the atmosphere surrounding Earth. They are rare events, and the higher their energy, the rarer they get. The greatest mystery surrounding them is their origin, but also their nature remains contentious. High statistics and high quality data are needed to make progress on both fronts, which means scanning the largest possible volume of atmosphere for EAS. The current ground based experiments run up against natural boundaries limiting their expansion and present difficulties when comparing data obtained in northern and southern latitudes, as the fraction of the sky common to both experiments is limited. Therefore, space based instruments observing the atmosphere from above with full-sky coverage have long been considered the logical next step in the evolution of UHECR experiments [1].

The International Space Station (ISS) with its existing infrastructure and support systems is a natural first step on this way into space, and JEM-EUSO [2] is a scientific mission under development with the aim of identifying the astrophysical origin and nature of UHECRs from the ISS. JEM-EUSO detects UHECR induced EAS by looking down onto the earth atmosphere. It has a telescope with a large ($\pm 30^\circ$) Field of View (FoV) imaging the atmosphere below the ISS onto an array of UV sensitive Multi-Anode Photomultiplier Tubes (MAPMTs) [3]. The MAPMTs (Hamamatsu Photonics R11265-03-M64) have 8×8 pixels and for readout purposes 2×2 MAPMTs are grouped into one Elementary Cell (EC). The First Level Trigger (FLT), which is the subject of this article, works at the level of these ECs. Nine ECs form one Photo-Detector Module (PDM), which is the basic unit for the Second Level Trigger (SLT). The Focal Surface (FS) is organised in 137 PDMs. Together these PDMs cover the FS of the telescope with $\sim 3.2 \cdot 10^5$ MAPMT pixels. A detailed description of the electronics and data acquisition for JEM-EUSO can be found in [4], see also 1.

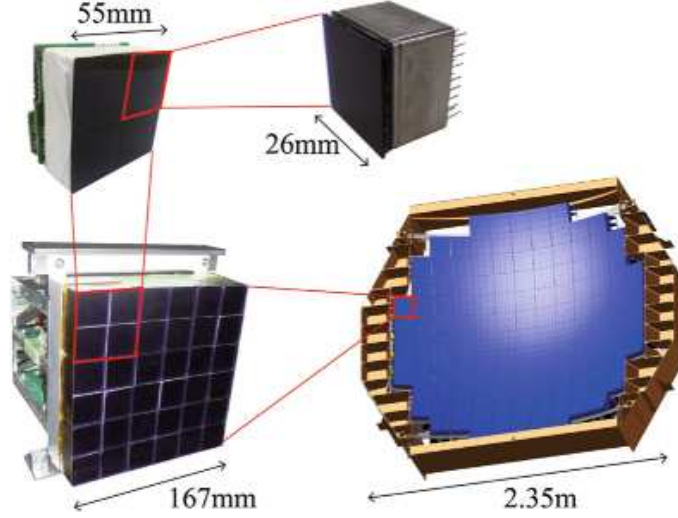


Figure 1: Structure of the Focal Surface. The 2.5 m surface is divided in 137 PDM modules. Each PDM is filled with 9 ECs, with 4 MAPMTs each. The bottom left corner shows the prototype of the mechanical structure with 36 MSPMTs installed. Figure taken from [4].

30 The observational concept of JEM-EUSO [5] is based on recording both
 31 the fluorescence light emitted during the evolution of EAS as well as the
 32 reflected Čerenkov light if the EAS's Čerenkov cone hits a reflective surface
 33 as it reaches the ground. EAS from the interaction of UHECRs or neutrinos
 34 in the atmosphere will - for 10^{20} eV EAS - typically result in a few thou-
 35 sand photons detected by the JEM-EUSO detector within a few hundred
 36 microseconds. Owing to the large FoV the expected rate of such ultra high
 37 energy EAS are approximately one per day. Depending on both the energy
 38 and the zenith angle of the EAS, its image may be contained inside a sin-
 39 gle EC or may cross a few PDMs as it is imaged onto across the FS. EAS
 40 develop within the lowest 15 km of the Earth atmosphere, so that their dis-
 41 tance to the ISS, which orbits earth at a height above ground of about 400
 42 km, can be considered unchanging whatever an individual EAS' zenith angle
 43 happens to be. With that EAS' angular speed across the FS to first order
 44 only depends on the EAS' propagation direction relative to the respective FS
 45 pixels' direction of view. As EAS traverse the atmosphere at essentially the
 46 speed of light, and from the height of the ISS a single square MAPMT pixel's
 47 FoV's diagonal measures roughly 750 m on the ground, it takes about $2.5 \mu\text{s}$
 48 for horizontal EAS' image to traverse the diagonal of a MAPMT pixel. As

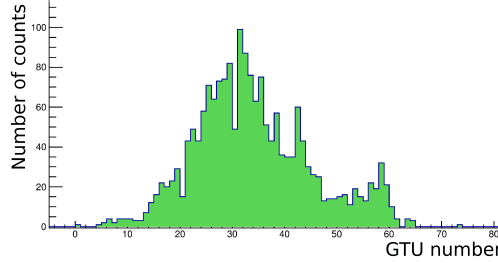


Figure 2: Light curve (number of photoelectron counts) integrated over the shower profile and plotted over time along the shower axis for a simulated UHECR EAS of 2×10^{20} eV (no background is added). Time is measured in GTUs (see Section 2).

bandwidth for data transmission from the ISS back to Earth is limited, 2.5 μ s, the so-called Gate Time Unit (GTU), was adopted as the basic unit for digitization at JEM-EUSO. Given the distance between EAS and ISS, JEM-EUSO must be able to detect single photons. The front-end electronics works in single photon-counting mode, which means that HV and electronics gain are adjusted such that after digitization one digital increment corresponds to one photoelectron (PE) count released from the MAPMT's photocathode.

In this paper we discuss the FLT algorithm specific to the identification of UHECR and neutrino induced EAS. Fig. 2 shows the temporal evolution in units of GTU of the MAPMT signal for typical simulated proton EAS of energy $2 \cdot 10^{20}$ eV, viewed by the JEM-EUSO telescope under an angle of 60° simulated with ESAF [6]. In Fig. 3, the top panel shows the spacial distribution of simulated EAS scintillation light emission projected back onto the Earth's surface for EAS with a common energy of $E = 10^{20}$ eV traversing the atmosphere under zenith angles of (a) $\theta = 30^\circ$, (b) $\theta = 60^\circ$ and (c) $\theta = 75^\circ$. The inset in the lower left corner of this panel puts these showers into the context of the FoV of the whole FS of JEM-EUSO. The bottom panel presents the image of the EAS in (b) as it would be seen by the JEM-EUSO telescope: the optics inverts the direction of motion, and the photon counts per pixel are integrated over the EAS duration ($\sim 100 \mu$ s).

Looking down from the ISS the FLT has to identify these events in the presence of various backgrounds: UV albedo, transient atmospheric phenomena, and artificial light sources in cities, along transportation networks, and on ships and airplanes. The ISS is moving at about 7.6 km/s so that stationary light sources on the ground stay within the FoV of a single pixel for about 70 ms. Such anthropogenic lights, as for example cities, are in the FoV

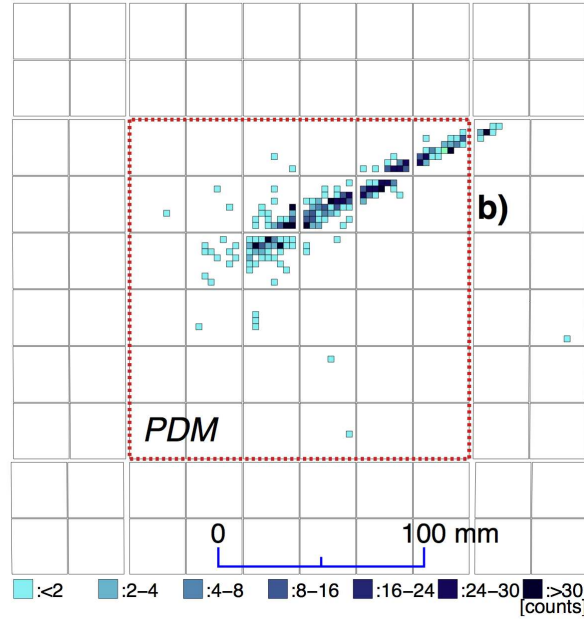
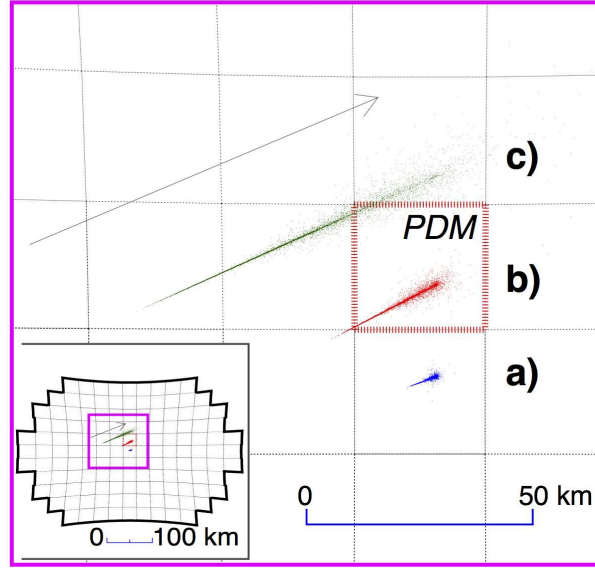


Figure 3: The top panel shows the 3 EAS with the same energy of $E = 10^{20}$ eV but impinging on the atmosphere under different zenith angles of (a) $\theta = 30^\circ$, (b) $\theta = 60^\circ$ and (c) $\theta = 75^\circ$. The arrow indicates the direction of the EAS transit on the FS. The inset on the bottom left and the grey grid shows how the FoV is imaged on the telescope's FS. The distance scale on this panel refers to the distance the shower develops over as projected onto the Earth surface. The bottom panel shows the image (inverted over as projected onto the Earth surface) on shower (b) as recorded (integrated over time) by the JEM-EUSO telescope. The distance scale here refers to distance on the FS. The regions enclosed by thick dashed lines in both panels refer to the same PDM. Image taken from [7]. UV background is not added in these plots.

on average during only $\sim 10\%$ of an ISS orbit [7]. Transient Luminous Events (TLEs) within the atmosphere, like electric discharges (Elves, Sprites, Blue Jets and lightning) as well as meteors will have their own triggering schemes to support a separate science program with JEM-EUSO and will be suppressed by the UHECR and neutrino oriented FLT on the basis of time and light intensity structures. The expected rate of TLEs is $\sim 700/\text{day}$ [8]. The greatest uncertainty is associated with the very slowly varying backgrounds associated with the albedo of the atmosphere: its scattering and reflection of starlight - light reflected from the moon and planets - and airglow. On a clear night the resulting diffuse flux of such photons is expected to amount to between 300 and 1000 photons $\text{m}^{-2} \text{ns}^{-1} \text{sr}^{-1}$, an expectation derived from existing measurements [9, 10].

In this paper we show the results of the offline testing of the current JEM-EUSO FLT concept using data from three main sources rather than simulation: data collected during a high altitude balloon flight over Canada (EUSO-Balloon), laser shots recorded in coincidence with a ground based JEM-EUSO prototype (EUSO-TA), and data taken in a laboratory setup where realistic background scenarios can be explored (EUSO@TurLab).

2. Technical requirements

As mentioned before, working on the ISS imposes severe bandwidth constraints on data transfer to the ground. On top of that there is a $\sim 1 \text{ kW}$ limit on power consumption for the whole telescope, including the readout and trigger electronics, high voltage for the MAPMTs, and monitoring. This constraint for example means that triggering cannot be substituted for by massive only computing. Data rate considerations also played into choosing a GTU of $2.5 \mu\text{s}$ for digitization in time [and the number of 128 GTUs per event](#), meaning that each event record will contain the timing evolution of a signal over a time span of $320 \mu\text{s}$.

Since the ISS is far above the parts of the atmosphere where EAS develop, the MAPMTs have to be able to detect faint signals using photon-counting. Therefore, 8 bit full scale per pixel is sufficiently large. Under these conditions the total data rate from the telescope before the FLT would be of order $3.2 \cdot 10^5 \text{ pixel/FS} \times 4 \cdot 10^5 \text{ GTU/s} \times 8 \text{ bit/pixel} \approx 1 \text{ Tbps}$. To achieve the required overall data reduction of $\sim 3 \cdot 10^6$ the FLT will have to reduce the trigger rate to $\sim 1 \text{ Hz/EC}$, and the PDM based second level trigger to $\sim 0.1 \text{ Hz/FS}$ [11, 12].

111 In the following section, specifics of the FLT logic aimed at detecting
 112 UHECR and neutrino induced EAS, which are the main scientific objective
 113 of the JEM-EUSO mission, are considered. While TLEs and other bright but
 114 slow atmospheric events are part of the exploratory objectives of the mission,
 115 we will not consider them here.

116 3. The 1st Trigger Level: Persistency Tracking Trigger

117 Persistency is a measure of how long in time a signal “persists” or stays
 118 within the FoV of a particular MAPMT pixel. As detailed above a GTU
 119 roughly reflects the time horizontal EAS right under the ISS need to travel
 120 the diagonal of a MAPMT pixel. To cross the FoV of an EC, where the FLT
 121 operates, takes up to 45 μ s for EAS, milliseconds for lightning, hundreds of
 122 milliseconds for meteors, and seconds for cities or airplanes. These differences
 123 in persistency and the fact that the signal moves from pixel to pixel as EAS
 124 pass through the FoV of the telescope were exploited in designing the FLT
 125 logic, which is described in [13]. Here we give a summary of how an EAS
 126 trigger is formed at the EC/FLT level.

127 Unless EAS develop along the line of sight of the telescope, the image of
 128 its fluorescence trail in the atmosphere can be tracked across the FS. Tracking
 129 discriminates EAS images against accidental coincidences of background light
 130 fluctuations.

131 To decide if a single MAPMT pixel is seeing signal above a slowly vary-
 132 ing, non-negligible background, this background level has to be estimated.
 133 Two different approaches were pursued: The pixel-based estimate sets one
 134 threshold for a whole MAPMT. To obtain this threshold the average over
 135 a 128 GTU data packet is calculated for each pixel in the MAPMT, and
 136 the maximum of these 64 averages becomes the background estimate and
 137 threshold for triggering in the next 128 GTU packet. Stationary or slowly
 138 moving anthropogenic light sources within the atmosphere are automatically
 139 suppressed by this method. The group-based estimate divides the whole EC
 140 into 32 groups of 2×4 pixels, calculates the 128 GTU averages per group,
 141 and chooses the maximum of those 32 averages as the threshold for all pixels
 142 in the EC during the next 128 GTU packet. The threshold here is a digital
 143 value as all calculations are done after digitization. No analog thresholds
 144 are used. The results presented in this paper are based on the pixel-based
 145 estimate which turned out to be better performing overall.

146 For tracking purposes each MAPMT's pixels are grouped into square 3×3
 147 cells. Each pixel belongs to more than one cell, but since cells do not span
 148 MAPMT borders, pixels along the edges and in particular at the corners of
 149 each MAPMT belong to fewer cells than the central pixels. The first panel
 150 in Fig. 4 shows three such overlapping cells on a single MAPMT's pixel grid
 151 outlined in dark orange. Since it includes neighboring MAPMTs' pixels, the
 152 3×3 pixel patch surrounded by the dashed grey line on the other hand does
 153 not constitute a cell. A total of 36 cells exists within each MAPMT. Each
 154 pixel for which its digitized signal in a certain GTU of a 128 GTU packet
 155 surpasses the threshold value $n_{\text{thr}}^{\text{pix}}$ for that MAPMT (determined from the
 156 data in the preceding 128 GTU) contributes to each of the cells it participates
 157 in. Apart from the pixel level signal threshold there also is a cell level signal
 158 threshold $n_{\text{thr}}^{\text{cell}}$. For a typical background level of one PE per GTU per pixel
 159 these thresholds $n_{\text{thr}}^{\text{pix}}$ and $n_{\text{thr}}^{\text{cell}}$ would normally be set to 3 and 31, respectively.

160 Persistency at the pixel level is evaluated based on two more paramete-
 161 rs that unlike those introduced above do not depend on the background
 162 situation: a pre-determined range of consecutive GTUs N_{pst} over which per-
 163 sistency is to be evaluated, and a limit N_{ctd} on the number of GTUs within
 164 that range for which pixels in the cell are above threshold. Standard values
 165 for the pixel related parameters N_{pst} and N_{ctd} are 5 and 3 GTUs, respec-
 166 tively. These values were determined by means of simulations of EAS signals
 167 and preliminary tests on MAPMT fluctuations with the aim of keeping EAS
 168 signals and rejecting background fluctuations. Persistency of a signal at the
 169 EC level is also monitored and similarly checked by two parameters for a
 170 maximum allowed number of GTUs above threshold $N_{\text{GTU}}^{\text{thr}}$ in a GTU range
 171 N_{GTU} . Too many GTUs with signal indicate high persistency, which is the
 172 hallmark of non-EAS induced events like lightning or meteors. This GTU
 173 range is started at the GTU in which for the first time a cell threshold is
 174 surpassed. Typical values for EAS identification are $N_{\text{GTU}} = 73$ and $N_{\text{GTU}}^{\text{thr}} =$
 175 72. These two values were decided according to the following considerations.
 176 N_{GTU} is determined by the number of GTUs remaining after the first trigger
 177 till the end of the packet. For technical reasons it was decided to have the
 178 trigger at GTU 55 of a packet. The $N_{\text{GTU}}^{\text{thr}}$ indicates that all the events are
 179 accepted unless the PDM continues triggering for every GTU after the first
 180 trigger till the end of the packet. Both values could be fine tuned in future, if
 181 needed. In particular $N_{\text{GTU}}^{\text{thr}}$ could be easily shortened by a few GTUs without
 182 impacting the trigger efficiency on EAS.

183 As mentioned above power consumption is a major constraint on the ISS.

184 The current implementation of the FLT was programmed into and tested
 185 on a Xilinx Virtex6 model XC6VLX240T [14] FPGA. Given that it required
 186 only $\sim 7\%$ of this FPGA's resources to accommodate the logic needed for 1
 187 EC, it is expected that one such FPGA can host all 9 ECs belonging to one
 188 PDM. The graphic in the following figure reflects the FPGA architecture in
 189 its reference to the FPGA's various adders.

190 Fig. 4 uses an event recorded by EUSO@TurLab ² to illustrate how a
 191 FLT is formed. The background estimate was derived from the the preceding
 192 GTU packet as $n_{\text{thr}}^{\text{pix}} = 5$ and $n_{\text{thr}}^{\text{cell}} = 65$; the average background light level
 193 for that packet was of order 3.4 PE/GTU/pixel. Using the standard values
 194 of $N_{\text{ctd}} = 3$ and $N_{\text{pst}} = 5$, the graphic follows the cell's pixels' recorded PE
 195 counts for an EAS-like event created by a line of LEDs mimicking an almost
 196 vertical shower mostly staying in the FoV of the cell's central pixel. In the
 197 first panel of Fig. 4 the cell is highlighted in dark orange on the MAPMT's
 198 pixel map. The next five sub-panels of the figure after the MAPMT pixel
 199 overview show two pixel maps each for that same cell. The five sub-panels
 200 represent the five successive GTUs following the cell's first threshold crossing.
 201 The pixel map on the left in each sub-panel shows the raw PE counts recorded
 202 per MAPMT pixel in the respective GTU. Using the estimated $n_{\text{thr}}^{\text{pix}} = 5$
 203 PE background as threshold then leads to the pattern of threshold-crossing
 204 pixels displayed on the right of the sub-panel with each pixel's background
 205 subtracted PE signal estimates. The sum of that signal above background
 206 is then compared to the cell threshold $n_{\text{thr}}^{\text{cell}} = 65$ PE. In summary: for the
 207 chosen cell and its five GTUs after the cell's threshold crossing, at least one
 208 pixel in the cell crosses the pixel threshold for each GTU, and the total signal
 209 strength accumulated within this cell in each GTU is enough to contribute
 210 to the EC wide evaluation of the event. Therefore, at GTU step 5 the
 211 corresponding adder (T) is incremented. The EC-wide check with regard to
 212 the GTUs during which the signal passes through all the other EC cells is
 213 summarized in the last panel, where the content of the adder T is finally
 214 checked before a FLT is issued (or not) to the PDM for second level trigger
 215 purposes.

216 Persistence is the main concept behind the FLT implementation. In the
 217 following section we will discuss how this current implementation performs
 218 in the presence of background, using data recorded with EC modules in dedi-

²see Section 4.1 and Fig. 11 for more details.

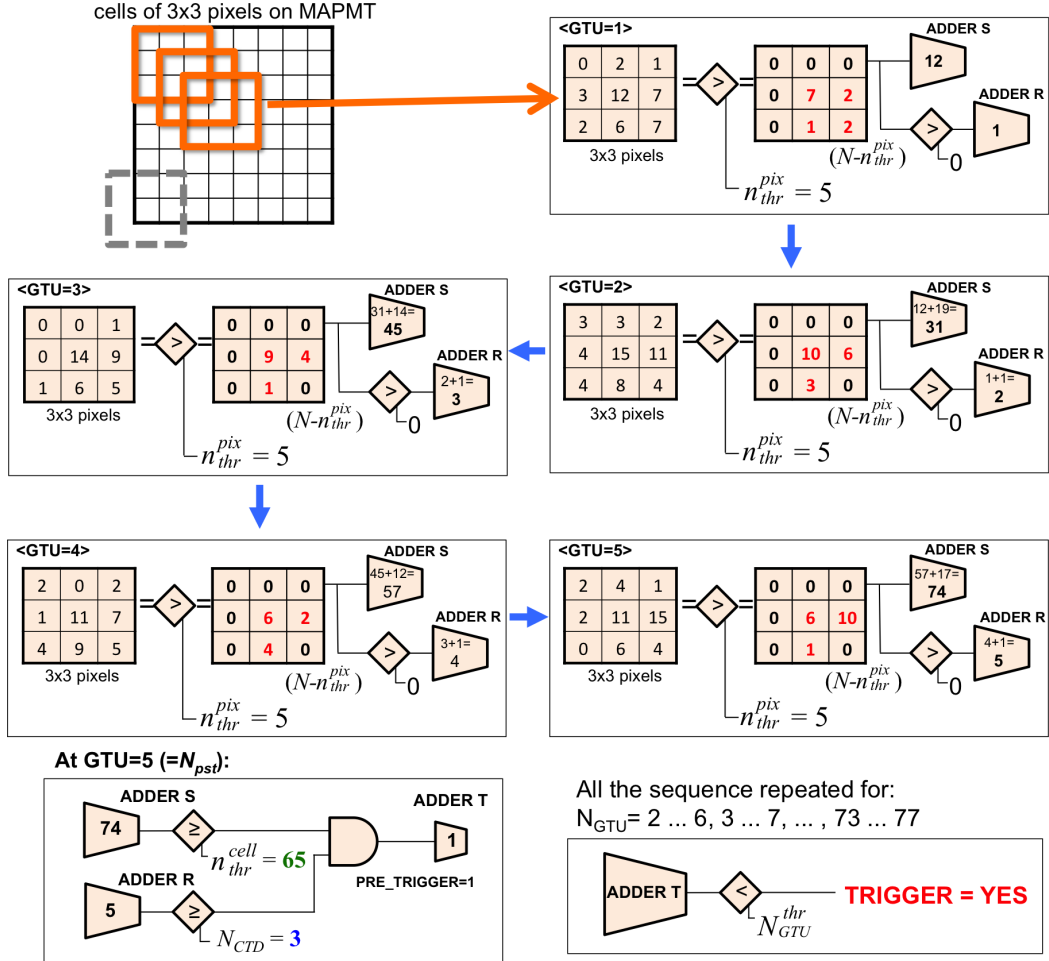


Figure 4: The FLT implementation at the level of the 3x3 cells. See text for details.

219 cated experiments reflecting different aspects of space based EAS observation
220 in the Earth atmosphere.

221 4. FLT Tests using experimental data

222 The trigger efficiency as a function of EAS energy (commonly referred
223 to as the trigger efficiency curve) captures an important aspect of the ex-
224 periment’s sensitivity. Several publications already discussed the expected
225 JEM-EUSO sensitivity given the current FLT implementation, but so far
226 they were purely based on MC simulations [5, 7, 15]. These simulations all
227 assumed Poisson fluctuations on a UV background intensity that is constant
228 and uniform across the FS.

229 In this section we report on tests performed offline using data taken with
230 actual ECs in three very different environments, each addressing specific chal-
231 lenges JEM-EUSO is expected to face during observation from the ISS: data
232 collected by the EUSO-Balloon flight in 2014 [16], measurements performed
233 by EUSO@TurLab at TurLab [17], and observations in coincidence with a
234 Telescope Array (TA) air fluorescence detector by EUSO-TA [18]. These
235 data sets allow to test the trigger system in very different and complemen-
236 tary ways. EUSO@TurLab provides the possibility to control lighting and
237 create realistic event patterns and persistencies, EUSO-Balloon takes data
238 under space-like conditions, and EUSO-TA allows comparison with a well
239 calibrated existing ground-based fluorescence detector.

240 4.1. Tests with TurLab measurements

241 The two main aspects of the FLT that were tested at TurLab, [located at](#)
242 [the Physics Department of the University of Turin \(Italy\)](#), were the adequacy
243 of its background estimation and the ability to trigger on EAS while sup-
244 pressing other signatures, [such as cities, meteors, lightnings, discontinuities](#)
245 [in the luminosity due to the presence of clouds, variation in soil condition,](#)
246 [moon phase, etc.](#) All these phenomena have variable intensity, duration and
247 extension. Table 1 gives typical values expected for JEM-EUSO for a subset
248 of these conditions which have been reproduced at TurLab to test the trigger
249 logic.

250 Being 15m under ground, the ambient light level in the TurLab laboratory
251 [17] is several orders of magnitude lower than that of the darkest night sky.
252 Using artificial light sources therefore puts the ambient light levels as well as
253 the distribution of light in the lab under the control of researchers.

light source	intensity (cts/pix/GTU)	duration	extension	variability
UV glow	0.5 - 5	orbit	EC	water, soil, cloud
Urban	3 - 30	seconds	pixel - EC	village - city
EAS	3 - 30	$\sim 100\mu s$	PMT (track)	EAS energy
Meteor	3 - 100	seconds	EC (track)	magnitude

Table 1: Variability of the signal expected for JEM-EUSO due to different light sources in the FoV of the telescope, ranging from steady UV nightglow to localized and impulsive light bursts such as cities, EAS, meteors.

At TurLab a rotating tank of 5 m diameter provides the stage on which light emitting as well as light reflecting installations are made. EUSO@TurLab consists of one EC hung off-center from the ceiling above this rotating tank. While in principle the EC can be moved radially it was kept at a radius of roughly 2 m from the center of the tank. The optics imaged 1 cm² on the tank's surface onto one pixel 2 m above the tank surface, giving it a FoV of the order of 10⁻⁵ sr, which is only one order of magnitude larger than a JEM-EUSO's pixel. This means that if the adjustable speed of the tank rotation were to be around two minutes, the time it takes a stationary source on the tank surface to cross a pixel would be the same as it will be for JEM-EUSO looking down on Earth from the ISS.

As outlined before, the data acquisition (DAQ) in JEM-EUSO will be a seamless sequence of 128 GTU long packets. At EUSO@TurLab the EC's ASIC is read out by a test board which transfers the data to a PC, and this system both limits a data packet to 100 GTUs and imposes a 50 ms deadline between two consecutive packet acquisitions. In other words, at EUSO@TurLab 100 GTUs = 250 μs of data are taken every 50 ms, and a stationary light source on the tank surface would have moved through 50% of the FoV of a pixel during that deadline if the tank rotated with a period of 2 min. It was not possible to synchronize the DAQ with the tank rotation, [because it was not foreseen by the hardware setup. Naturally, the synchronization would have allowed to determine exactly the location in the tank responsible for each trigger.](#)

Fig. 5 shows the various components of the TurLab setup. It shows the EC suspended from the ceiling, and various installations on the tank surface designed to emit or reflect light in ways that mimic both anthropogenic and natural lighting situations as they would be seen by JEM-EUSO from the

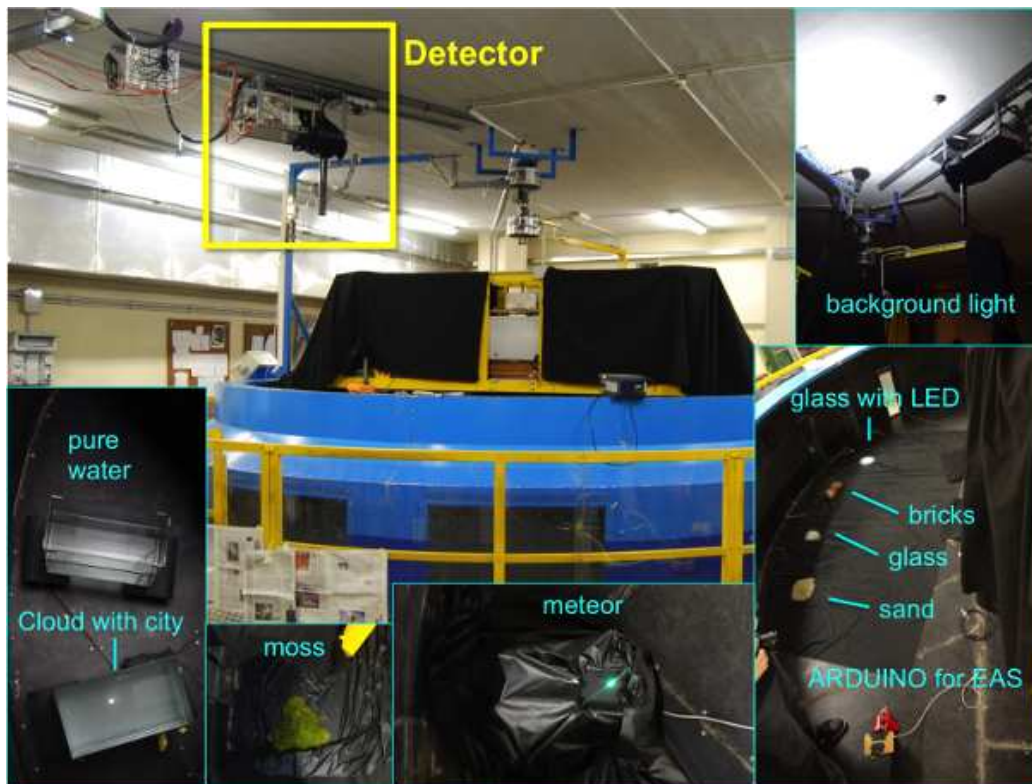


Figure 5: The TurLab rotating tank. The black tube on the ceiling shows the collimator of the experimental setup used to mimic the JEM-EUSO telescope. Light sources and materials used to mimic other phenomena are shown as well.

281 ISS. Light scattered or reflected from the atmosphere back into space - the
 282 Earth's [albedo](#) - depends on atmospheric conditions like e.g. the presence or
 283 absence of clouds, the reflectivity of the Earth's local surface, and lighting
 284 conditions like the phase of the moon³ or the presence and density of human
 285 habitation.

286 In Fig. 5 one can also see how sand, moss, ground glass, pure water,
 287 and a brick were used to mimic the reflection of night-sky light from soils,
 288 forests, snow, water, and rocky surfaces, respectively. Water clouded by dis-
 289 solved particles and illuminated from below is used to mimic clouds, and if
 290 illuminated from below cloud cover over e.g. a city. An oscilloscope screen
 291 displaying Lissajous traces mimics meteor tracks. As lighting can be con-
 292 trolled, the TurLab tank allows to verify the performance of the background
 293 estimation under realistically varying lighting conditions. Fig. 6 and Fig. 7
 294 show two examples of recordings of such features with EUSO@TurLab.

295 In Fig. 6 LED light reflected from ground glass is used to simulate the
 296 distributed individual light sources of a city. Each of the four frames in each
 297 sub-panel refers to one of the four MAPMTs in the EC. The upper panels
 298 show the respective MAPMT's integrated PE counts, in other words the
 299 light curve of a city passing through the FoV of the EC. The city entering
 300 and exiting the MAPMT's FoV as time progresses is clearly visible in each
 301 MAPMT's light curve. The lower row shows 2D pixel maps for the EC's
 302 MAPMTs, with the PE counts per pixel for just one GTU on the left, 10
 303 GTUs in the center, and 100 GTUs on the right. The red lines in the light
 304 curves show the range of GTUs that are used, with the single GTU pixel maps
 305 being the first GTU under both red lines; the start times of this integration
 306 is the same for all three ranges. Fig. 7 shows the data recorded while passing
 307 over the oscilloscope repeating a straight line Lissajous figure taking about
 308 one second to complete. While the complete picture emerges after integrating
 309 over 1500 GTUs (right panel), the signal still is contained in a single pixel
 310 when integrating over only 10 GTUs (center panel).

311 An Arduino board [19] controlling a line of 10 white LEDs was used to
 312 mimic a single EAS propagating through the atmosphere at the speed of
 313 light, resulting in a total duration of about 40 GTUs. As can be seen in
 314 Fig. 8 this signal does no longer stay within one pixel during 10 GTUs, with
 315 the center of light moving visibly between subsequent GTUs.

³Fluorescence observation of EAS is not possible during daytime.

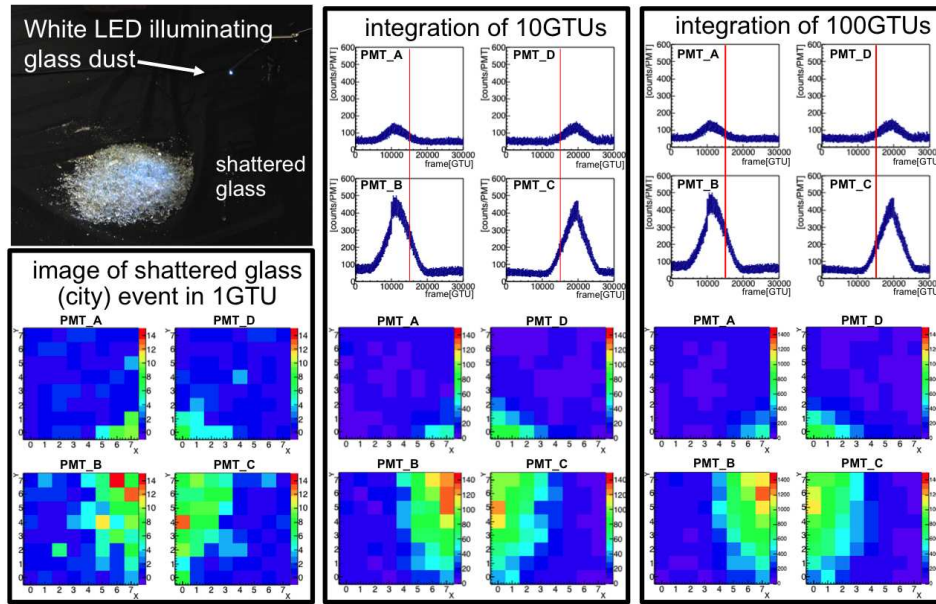


Figure 6: Reproduction of an extended light source, like in case of a city. Top-left plot is a picture of the shattered glass; bottom left plot is the image detected by the MAPMTs in 1 GTU. The right-top plots shows the temporal evolution of the same scene with different time integrations (10 and 100 GTUs). The bottom plot shows one frame per integration taken at the time indicated by the red line in the above plots.

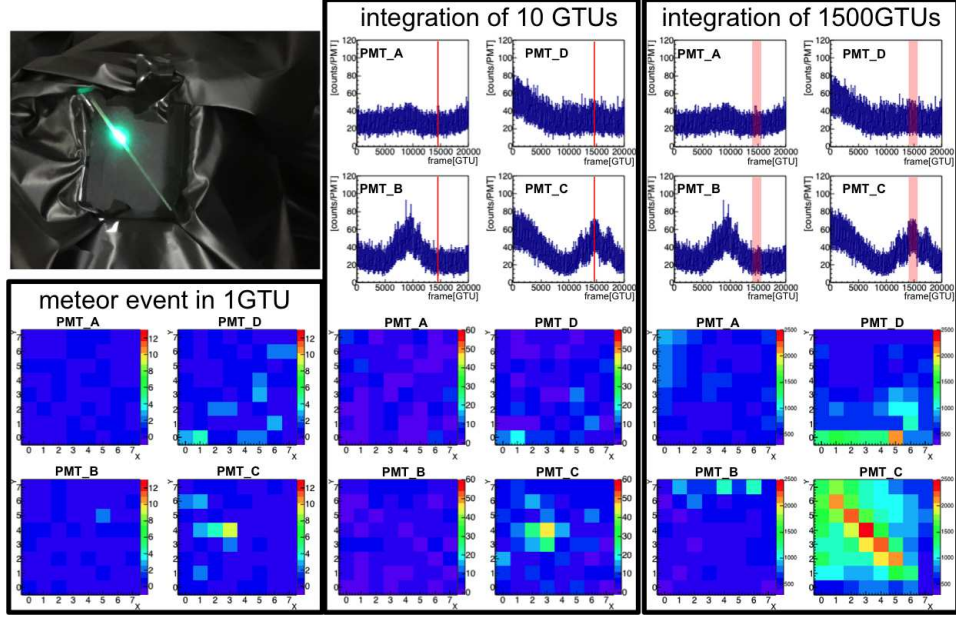


Figure 7: Reproduction of a meteor-like track. See Fig. 6 for details about the meaning of each plot in the figure.

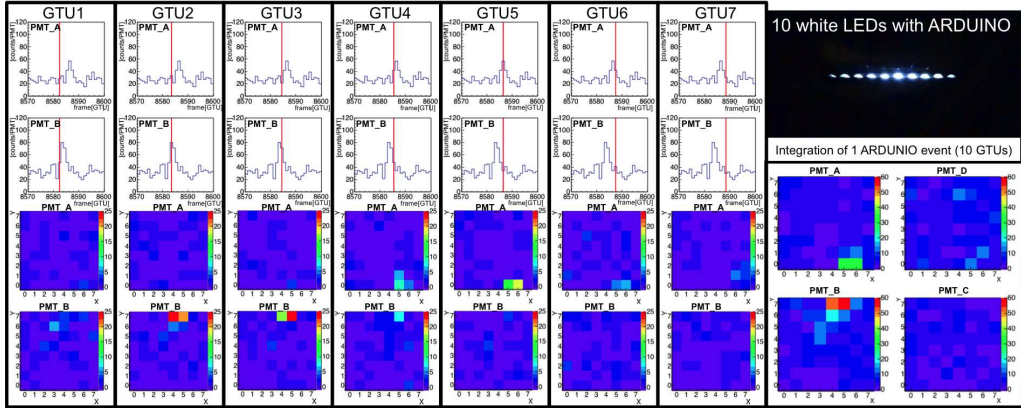


Figure 8: Reproduction of a cosmic ray-like track. Top-right picture shows the integrated light sequence reproducing a cosmic-ray track. Bottom-right plot shows the integrated number of counts during the light sequence. The left part of the figure displays 7 frames of 1 GTU each taken during the reproduction of the track. The time at which the frames are taken is shown in the above corresponding plots which present the time evolution of the total number of counts recorded by the MAPMT.

316 The FLT estimates the background for the current data packet from the
 317 data collected in the preceding data packet. To mitigate possible adverse
 318 effects of the DAQ-imposed deadtime between the acquisition of consecutive
 319 packets on background estimation the tank rotation was slowed to complete
 320 one rotation in 9 min, reducing the offset of a stationary light source between
 321 consecutive acquisition packets from 50% of a pixel to roughly 10% of a pixel;
 322 with the JEM-EUSO DAQ a 128 GTU offset at ISS speed would correspond
 323 to 0.5% of the pixel size projected onto the ground. Given the deadtime
 324 between the 100 GTU acquisition packets a total of ~ 3 seconds of data is
 325 collected during one 9 minute rotation.

326 The DAQ at TurLab collects that data “as is”: it simply reads out the
 327 PE counts for each MAPMT pixel in each GTU from the EC’s ASIC and
 328 writes them to disk. The subsequent trigger simulation is then implemented
 329 in VHDL⁴ according to the schema described in [13].

330 The light collected on one of the EC’s MAPMTs during a complete 9
 331 minute rotation of the tank is shown in the upper panel of Fig. 9. Changes
 332 in the background light level are clearly seen, and the various contraptions
 333 that precipitate them are labeled in the figure. The two bridges refer to the
 334 footbridges to cross the tank which, despite being covered by some black
 335 fabric, are a source of quite variable light reflection. In general the black
 336 fabric was used to make as dark as possible specific regions of the tank to
 337 help increase the dynamic range of the light intensity seen by the MAPMTs
 338 during the tank rotation. Pure water in a little transparent tank was used
 339 to mimic a mirror-like condition which induces much higher reflection. The
 340 yellow bar is a pole on the rotating tank which passes a few cm below the
 341 collimator, thus filling a significant portion of the FoV of the detector for a
 342 short time. The second panel of Fig. 9 shows the average PE count in the
 343 preceding 100 GTU packet for the pixel with the maximal average count in
 344 that same MAPMT, which is the value used for the threshold setting in the
 345 current packet. The final panel shows how the trigger simulation reacted
 346 to this input. It shows when FLT’s were issued based on signals in that
 347 MAPMT. Almost all triggers coincide with passing over the Arduino driven
 348 LED chain as it should be; the one that is not is due to a specific location
 349 near one of the two bridges crossing the tank where the variations of light
 350 reflection were still too fast to be compensated by the slower rotation of the

⁴Very high speed integrated circuit Hardware Description Language

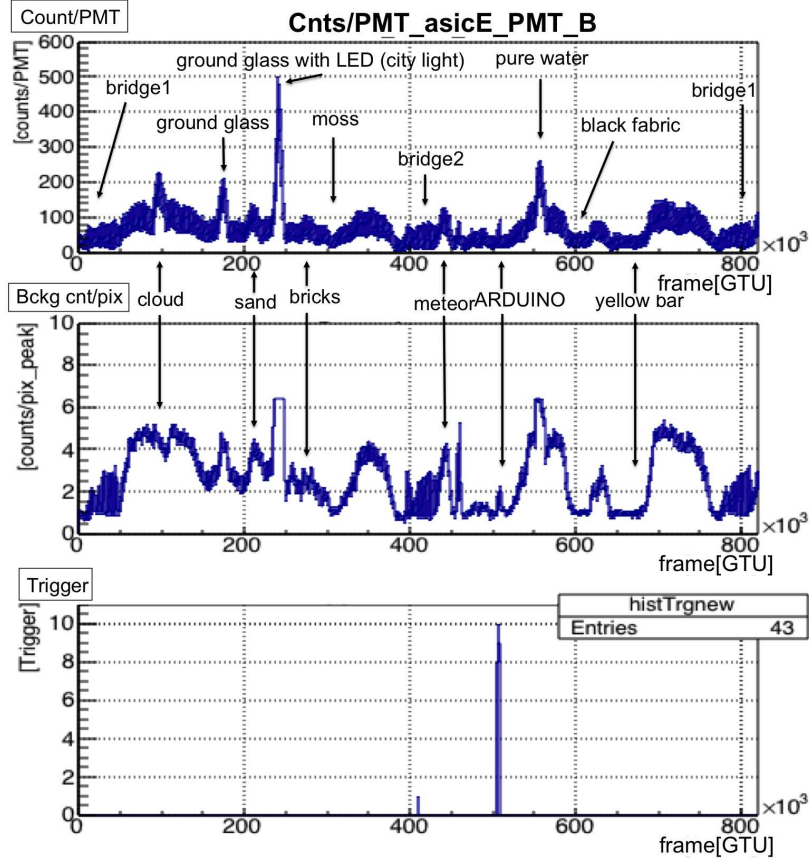


Figure 9: Reproduction of a full TurLab rotation with many types of light. Top plot shows the sum of the light collected by 1 MAPMT as a function of time. Middle plot shows the light intensity monitored by the pixel responsible to set the trigger thresholds of the MAPMT. Bottom plots show the triggered events. Except for two spurious cases [due to quite variable background conditions \(see middle panel\)](#) which could not be properly followed with the 50 ms dead time between packets, all the triggers coincide with the cosmic ray-like events generated by Arduino.

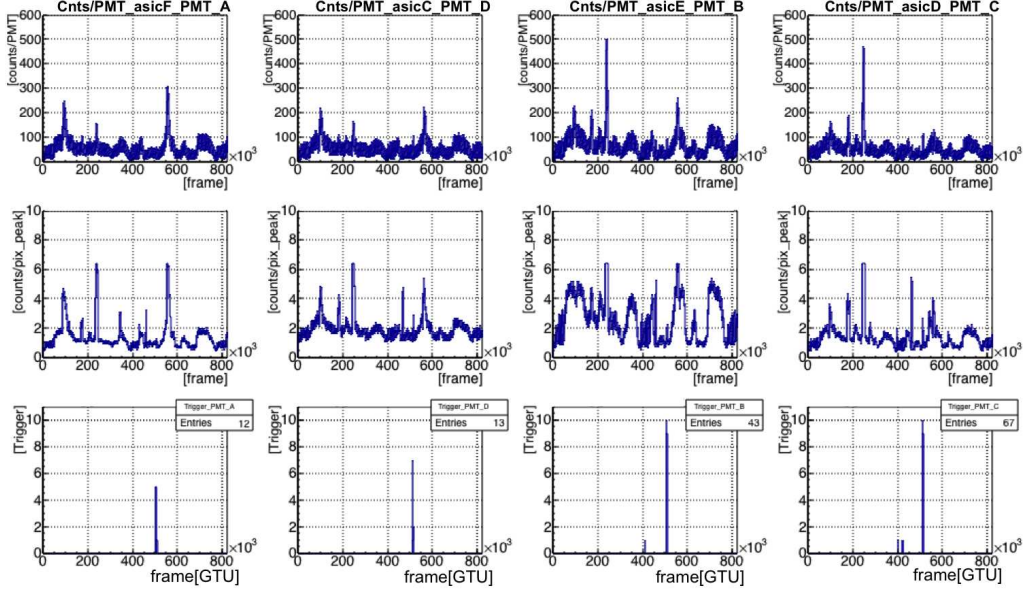


Figure 10: Reproduction of a full TurLab rotation with many types of light sources as shown in fig. 9 for the entire EC.

tank. Fig. 10 shows top and bottom plots of Fig. 9 for all four MAPMTs. A similar response is obtained for all four MAPMTs all along the rotation.

To directly assess the impact of background on the FLT trigger scheme for EAS dedicated measurements were made at TurLab with the tank rotation stopped and the EC stationary above the Arduino driven white LED strip simulating EAS. Ambient light levels then controlled the background to the LED induced signal. These ambient light levels were varied between 0.1 and 2.0 PE per pixel and per GTU, reflecting expectations for typical ISS observation background. The Arduino EAS were generated 1 ms apart in order to reduce the probability of recording such Arduino EAS in consecutive data packets, in which case the first EAS would set the background level for the second EAS. As the DAQ for EUSO@TurLab was not synchronized with track timing in the Arduino, extracting the packets containing a complete Arduino track required some event selection.

This selection started from a 4×4 pixel box in that MAPMT which contained the brightest part of the Arduino LED simulated EAS. The stationary tank was oriented such that the Arduino LEDs were all within the field of view of a single MAPMT and the Arduino EAS were always crossing the

369 same MAPMT pixels.

370 The LED sequence for these Arduino EAS was kept stable with ~ 30 PE
 371 at maximum, which corresponds to recording a $\sim 6 \times 10^{19}$ eV EAS in JEM-
 372 EUSO. A mask above the LEDs was used as an aperture to avoid unwanted
 373 reflections of LED light from nearby structures above the tank. The voltages
 374 supplied to the LEDs were also adjusted to dim the LEDs that were closer
 375 to the ends of the strip in an effort to provide a realistic EAS profile when
 376 the Arduino board sequentially lights up the LEDs in the strip.

377 If the PE count in the 4×4 pixel box smoothed over 5 GTUs exceeded
 378 the corresponding background estimate by more than 4σ , the data were con-
 379 sidered an Arduino EAS candidate. Such a candidate would subsequently
 380 be rejected if the excess occurred only in the first or last five GTUs of a 100
 381 GTU data packet, or if the preceding data packet also contained an Arduino
 EAS candidate. Fig. 11 shows PE counts for the relevant MAPMT over

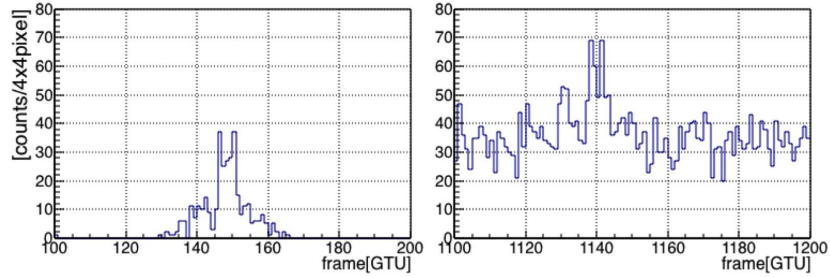


Figure 11: Example of light curves of two extracted Arduino events without background counts and with a background condition of ~ 2 counts GTU $^{-1}$ pixel $^{-1}$, respectively.

382 time in GTU units. The left panel shows a typical event produced by the
 383 LED strip without background. This highlights the event's original shape.
 384 The right one is a similar event produced under high background.

386 Events selected by this procedure were then fed into the VHDL trigger
 387 simulation. Table 2 shows that all the selected Arduino EAS also triggered
 388 in the FLT simulation. This together with the fact that when observing
 389 the rotating tank with its various implementations of atmospheric as well as
 390 ground-based light sources, albedo effects and generally varying background
 391 light levels gave rise to only a few spurious triggers under very specific con-
 392 ditions gives confidence that the current FLT implementation is ready for
 393 deployment in JEM-EUSO.

$N_{PEave./4\times4pix}$	$N_{triggered}$	$N_{extracted}$
0.21	24	24
0.36	13	13
0.56	15	15
0.78	21	21
1.00	26	26
1.23	16	16
1.47	26	26
1.69	22	22
1.93	20	20
2.11	31	31

Table 2: Number of triggered and extracted cosmic-ray-like-track events in various background photon level conditions. $N_{PEave./4\times4pix}$ indicates the average background level expressed in counts per pixel per GTU evaluated on a 4×4 pixel-box during the preceding packet of data, where no Arduino event was extracted.

394 4.2. Tests with EUSO-Balloon data

395 The EUSO-Balloon [20, 21] data taken during a 5 hour flight at 38 km
396 altitude in the vicinity of Timmins in Canada provides another testbed for
397 the FLT. Again the adequacy of the newly adopted background estimation
398 method with respect to keeping the trigger rate within the permissible bounds
399 in the presence of artificially and naturally encountered fluctuations in the
400 background lighting conditions as well as the FLT’s ability to trigger on
401 relevant optical phenomena was studied. While at TurLab the optics and
402 speed could be adjusted to match event duration and persistence in a pixel’s
403 FoV for the various phenomena recreated there, the EUSO-Balloon’s speed
404 and trajectory could not be controlled to that extent. On the other hand
405 EUSO-Balloon looked down on a real Earth environment just as JEM-EUSO
406 will, albeit from a much closer distance than the ISS. Thus were TurLab
407 strove to be realistic in an artificial environment, EUSO-Balloon was looking
408 at realistic settings compromising on perspective. Thus they each capture
409 different aspects of the challenges presented to an FLT operating at JEM-
410 EUSO.

411 In flight the EUSO-Balloon optics imaged a 60 km^2 surface area onto one
412 full PDM with its 9 ECs and a total of 36 MAPMTs. Just as at TurLab
413 the data acquisition did not allow to record data continuously, but took 128
414 GTU data packets at 18 Hz, translating into $320\text{ }\mu\text{s}$ of data recorded every

415 ~ 55 ms. Data taking on the balloon was organized into runs. A new run
 416 would be started after either 200 or 2000 data packets had been recorded.
 417 In total about $4 \cdot 10^7$ GTUs were recorded with the camera looking down on
 418 natural backgrounds like forests, lakes, and clouds, as well as city lights.
 Fig. 12 traces the time variation of a typical one of the 2304 pixels in the

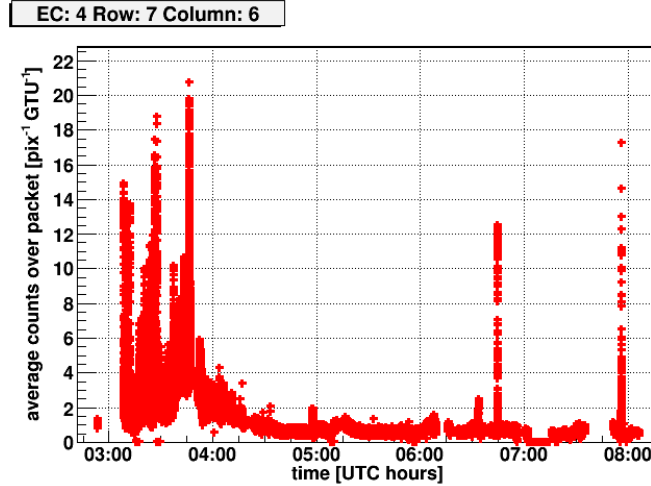


Figure 12: PE counts averaged over each recorded data packet for one typical pixel, covering the entire balloon flight.

419 camera. The city of Timmins for example crossed the field of view of that
 420 pixel between 03:00 and 04:00 UTC. The low counts at the center of the data
 421 taking period correspond to times when the balloon was passing over forests,
 422 lakes, and clouds. The highest PE counts were recorded when the balloon
 423 passed over an active mining operation.
 424

425 For about two hours a helicopter was flying under the balloon. The
 426 helicopter carried three different light sources: a UV LED, a Xe-Flashlamp
 427 and a UV laser. This sources had the purpose of calibrating the detector
 428 response and simulating EAS-like event patterns that in this paper are used
 429 offline to test the FLT. Data analysis focused on about one hour of balloon
 430 flight during which the helicopter was shooting underneath the balloon and
 431 in the FoV of the camera. During this hour no other system tests interfered
 432 with the measurements, and the flight path crossed over dark as well as bright
 433 areas. This is essential to test the FLT logic under extreme conditions.

434 The data was divided into two data blocks: One block with so-called
 435 “nominal” background levels as are expected for operation on the ISS with

436 ~ 0.6 PE/pixel/GTU averaged over a packet, and one block with a “high”
 437 background of 5–10 PE/pixel/GTU average over a packet in some parts of
 438 the FoV. This latter block of data was collected while passing over the mine
 439 or a city. About 15% of MAPMT pixels that were not working properly were
 440 excluded from the analysis.

441 The light sources on the helicopter were set up to emit signature patterns
 442 that each served a distinct purpose [22]. First in the sequence was a UV-LED
 443 (375 nm wavelength), the light output of which steadily increased with time
 444 over 12 GTUs. From the balloon this signal appears as a stationary source
 445 typically contained in a single MAPMT pixel. The UV-LED light output
 446 was kept stable throughout the night and designed to raise the signal level
 447 from ~ 1 to ~ 50 PE over the 12 GTUs in that pixel. This signal provides
 448 a normalization for the distance between helicopter and EUSO-Balloon and
 449 allows to determine an effective threshold for the FLT.

450 Next in the sequence was a laser pulse shot horizontally away from the
 451 helicopter. This laser shot was fired about 25 GTU after the end of the
 452 UV-LED signal, delivering ~ 5 mJ over 7.5 ns at a wavelength of 355 nm.
 453 Depending on where in the balloon’s FoV the helicopter happened to be
 454 at that time, it could take a maximum of 10 GTUs before the laser pulse
 455 would leave the balloon’s FoV. The number of photons scattered out of such
 456 a laser pulse roughly corresponds to the fluorescence light emitted at shower
 457 maximum from a $\sim 10^{20}$ eV EAS [according to ESAF simulations](#).

458 The balloon’s altitude being low compared to the ISS however meant that
 459 the ~ 400 m \times 400 m of a 3×3 pixel cell on the ground was crossed by the
 460 laser pulse in ~ 1 GTU, while the FLT is integrating over 5 GTU to establish
 461 a threshold crossing. To retain the ability for the FLT to trigger, the last
 462 light source in the sequence of light sources operated on the helicopter is a
 463 xenon (Xe) discharge lamp (wavelength 337 nm) emitting its light over ~ 8
 464 GTU, hereafter referred to as Xe flasher. This Xe flasher is triggered ~ 5 μ s
 465 after the laser shot, and its light curve reaches its maximum three GTU from
 466 its start, decreasing thereafter. This is reminiscent of the light curve along
 467 EAS, and four different flash intensities were used to mimic different EAS
 468 energies. A total helicopter light sequence therefore extends over ~ 50 GTU.

469 Fig. 13 shows the integrated PE count per pixel for a 128 GTU data
 470 packet containing a whole helicopter light sequence. An offset was subtracted
 471 throughout to highlight the excess along the laser track. The helicopter with
 472 the UV-LED and the Xe flasher was in the pixel at $x=5$, $y=25$. The left panel
 473 of Fig. 14 shows the PE sum for each GTU in that same data packet for the

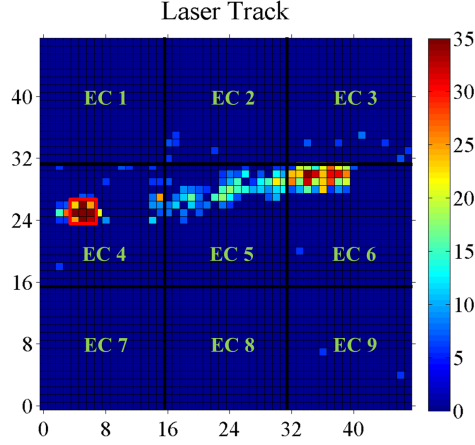


Figure 13: Image of a helicopter event obtained by integrating the counts in each pixel for the whole packet=1960 of RUN=043202 (128 GTUs). A threshold is applied to the minimum signal level to emphasise the location of the track. The UV-LED and Xe-flasher signals are centred around pixel at axis of abscissae $X=5$ and axis of ordinates $Y=25$.

474 3×3 pixel cell that is centered on the helicopter position at $x=5$, $y=25$. The
 475 UV-LED ramp can be seen to start from GTU 19. The laser shot is seen at
 476 GTU 55 and 56, and the Xe flash lamp is seen between GTU 58 and 65. The
 477 signal peaking at GTU 72 is attributed to an afterpulse in the Xe flasher.

478 Fig. 13 also identifies the ECs that the various signals appear in. For the
 479 UV-LED and the Xe flasher this is EC number 4, and the laser track starts
 480 in EC number 4 and then moves out through ECs 5 and 6. The right hand
 481 panel of Fig. 14 shows the number of ensuing FLT triggers per GTU for all
 482 ECs in that same data packet. The red line refers to triggers in EC number
 483 4, green to EC number 5, and blue to EC number 6. All signals are detected
 484 by the FLT algorithm, and the delay between the signal and trigger timings
 485 simply reflect the 5 GTU integration.

486 As the EUSO-Balloon DAQ and the helicopter light sequence each run
 487 on their own respective clocks and the DAQ recorded only $2.5 \times 128 = 320 \mu\text{s}$
 488 every $\sim 55 \text{ ms}$ ($\sim 0.6\%$), the vast majority of helicopter light sequences were
 489 not recorded. There will also be events where only a part of the helicopter
 490 light sequence overlapped with a DAQ data packet. Running the offline FLT
 491 simulation through the data, 274 events were found in which at least two
 492 ECs triggered the FLT algorithm.

493 Another peculiarity of EUSO-Balloon was that the optical module under

the balloon spun, constantly changing the FoV's alignment with respect to both surface features and laser direction, and did so at a varying rate. This meant that especially at the edges of the FoV stationary and other light sources would often enter or exit the FoV during the ~ 55 ms dead time between data packets. This complication should clearly be more relevant where stationary light sources on the ground play a significant role. As described above, the data was sorted into two blocks: one with nominal and the other with high background over locations lit up by human activity. The first block contains a total integrated time corresponding to ~ 8.5 seconds, and the latter corresponding to ~ 6.5 seconds, with about one order of magnitude more background light in this latter block's data on some parts of the FoV.

As expected the trigger rate under the more severe background conditions is higher: In the high background block of data the FLT algorithm triggered on 148 laser events and 59 others, while it triggered 126 laser and 17 other events in the nominal background data block. Assuming all other events are background, this puts the background rates for the current FLT trigger logic at 2.0 Hz per 9 ECs for the nominal background data block and 9.1 Hz per 9 ECs for the high background data block. Under both conditions the rate requirement of ~ 1 Hz per EC is met. In particular this means that despite the particular challenge posed by the combination of balloon spinning and DAQ deadtime the background estimation using the preceding data packet works well.

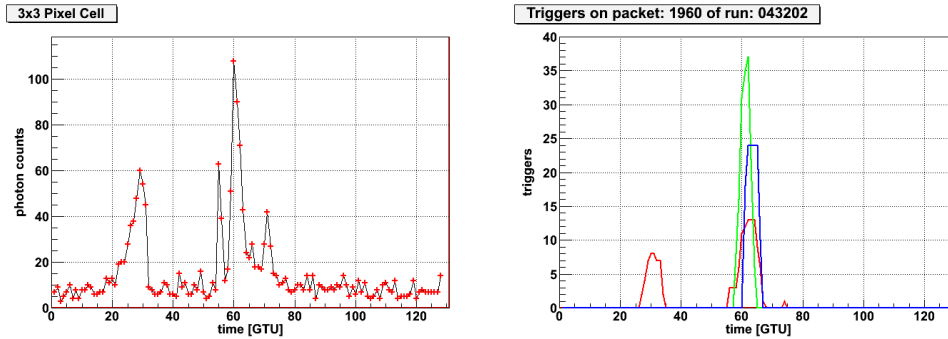


Figure 14: Left: Number of counts recorded in the 3×3 pixel-cell centred around ($X=5, Y=25$) during the entire packet. See text for details. Right: Sequence of trigger alerts in the different ECs crossed by the laser track during the entire packet. See text for details.

The event shown in Fig. 13 and Fig. 14 and triggered on by the FLT

517 algorithm can be used to estimate an energy threshold for EAS that would
 518 pass the FLT. Averaging over the seven lowest PE/GTU values that raise
 519 trigger alerts in EC number 4 the average signal excess becomes 81 ± 13 and
 520 the average background 39 ± 1 PE/GTU. This is a signal over background
 521 ratio of 2.1 ± 0.3 for the 3×3 pixel cell raising the trigger. Comparing
 522 this to ESAF simulations for EUSO-Balloon [23] under nominal background
 523 conditions this ratio is reached for vertical EAS initiated by a $\sim 5 \times 10^{18}$ eV
 524 proton. As the simulation also shows that showers at higher zenith angle
 525 provide higher signal/GTU, this value should be considered an upper limit
 526 for the energy threshold of FLT-triggered events recorded by EUSO-Balloon.
 527 Given its FoV and the measured CR rate at this energy this means that the
 528 FLT should trigger one event for every 24 hours of EUSO-Balloon livetime.

529 *4.3. Tests with EUSO-TA data*

530 TurLab measurements and EUSO-Balloon data were used to verify that
 531 the FLT and in particular its background estimation perform and meet the
 532 requirements under various realistic or even challenging background condi-
 533 tions. While comparing the lowest light level in EUSO-Balloon events that
 534 raised a FLT with simulation produced an estimate for the energy thresh-
 535 old in detecting cosmic ray particles, this is still a far cry from obtaining
 536 an efficiency curve for the FLT. This problem is addressed with data from
 537 EUSO-TA.

538 The EUSO-TA [18] telescope is a prototype of the JEM-EUSO space
 539 telescope with two 1 m^2 square Fresnel lenses. Just as for EUSO-Balloon its
 540 electronics comprise a full PDM with 9 ECs and 36 MAPMTs.

541 It is located right in front of the Black Rock Mesa (BRM) fluorescence
 542 detector (FD) site of the TA experiment in the Utah West Desert, USA [24].
 543 EUSO-TA's FoV of $11^\circ \times 11^\circ$ is contained within that of the BRM's FD and
 544 aligned such that it contains the vertical tracks from the pulsed 355 nm laser
 545 at TA's central laser facility (CLF). During TA data taking on the moonless
 546 parts of nights with amenable weather the CLF fires 300 vertical laser pulses
 547 of 3 mJ at 10 Hz every half hour. Providing atmospheric and calibration data
 548 for all three of TA's FD sites it is located centrally at an equal distance of
 549 21 km from each of the TA FDs, and therewith also 21 km from EUSO-TA.
 550 Depending on the offset between GTU boundaries and laser shot, laser tracks
 551 took 6 to 8 GTUs for their image to cross the PDM at EUSO-TA. The left
 552 panel of Fig. 15 shows an average over ~ 250 such CLF shots as recorded by
 553 EUSO-TA.

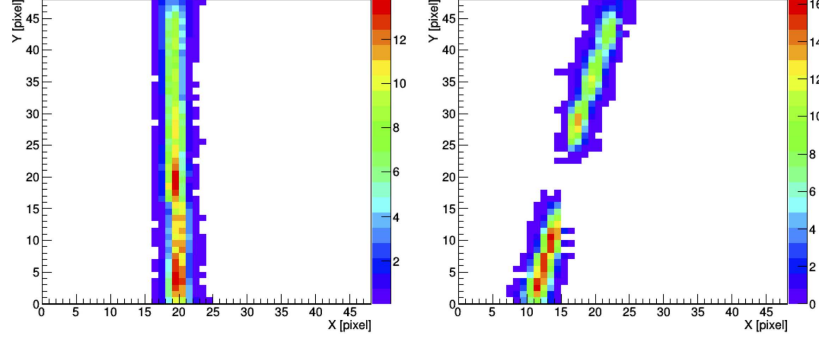


Figure 15: Left: an average of ~ 250 shots of CLF laser; right: an average of ~ 150 inclined shots of the Colorado School of Mines laser, located at 40 km from TA-EUSO, the missing part due to a non- functioning MAPMT in the center of the focal surface. The color scale on both pictures denotes the detector counts. Figure taken from [25].

554 The inclined laser track shown in the right panel of Fig. 15 is from a set
 555 of laser events produced with the help of a mobile UV laser belonging to the
 556 Colorado School of Mines. The missing piece in this laser track average was
 557 due to a defective MAPMT in EUSO-TA.

558 Also using a 355 nm laser the pulses from this mobile laser can be adjusted
 559 in intensity within a range of 1 to 86 mJ. As the laser itself is steerable, the
 560 geometry of the laser track can be varied more freely, and for the average
 561 over the ~ 150 laser pulses shown here the laser was shot at a distance of 40
 562 km with a pulse energy of 62 mJ.

563 Varying the laser pulse energies with this mobile laser at 34 km from
 564 EUSO-TA produced the trigger efficiency curve for the FLT that is shown
 565 in Fig. 16. As at these distances the laser pulses typically cross a few
 566 pixel/GTU, the FLT logic was adapted by setting $N_{pst} = 1$, while n_{thr}^{pix} and
 567 n_{thr}^{cell} were modified accordingly to keep the FLT trigger rates below the 1
 568 Hz/EC requirement. To determine the trigger efficiency, an external trig-
 569 ger, synchronized with the laser shooting, was supplied by the TA-FD to the
 570 EUSO-TA DAQ to always have the laser track inside a 128 GTU packet.
 571 The efficiency can then be determined by running the adapted FLT algo-
 572 rithm over these data packets and counting the packets that raise an FLT.
 573 Laser pulse energies between 3 and 5 mJ were used for the first four points
 574 with signal excesses above background < 50 PE over all pixels in that GTU.
 575 Above 50 PE overall signal excess, which corresponds to ~ 25 PE in the rel-

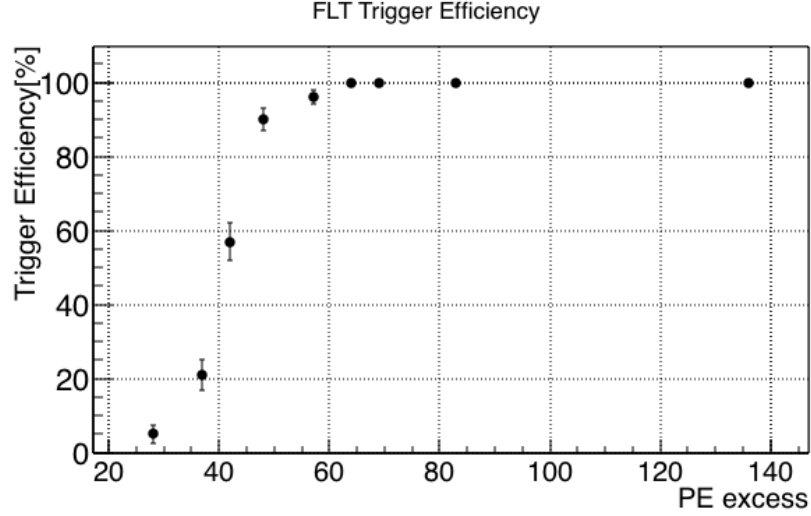


Figure 16: Trigger efficiency curve of FLT as a function of the signal excess recorded by TA-EUSO.

576 evant 3×3 pixel cell and 6 mJ pulse energy for this geometry, the trigger
 577 efficiency becomes 90% and higher. In a corresponding analysis for the 21 km
 578 CLF geometry the CLF's 3 mJ pulses were seen with 94% efficiency.

579 5. Conclusions

580 The FLT logic for use in JEM-EUSO as described above and implemented
 581 in VHDL was shown here a.) to work well in the presence of artificially
 582 produced as well as naturally encountered fluctuations in the background
 583 lighting conditions and b.) to keep the FLT rate within the permissible
 584 bounds while c.) being efficient at identifying event types with general EAS
 585 characteristics.

586 The FLT trigger as presented here is working at the MAPMT level and
 587 is based on the local persistency of a signal excess in a 3×3 pixel area, per-
 588 sisting a few GTUs. To achieve this an automatic evaluation of the average
 589 background level is derived from the preceding data package, as strategy that
 590 has proven successful even when individual data packages were separated by
 591 up to a few hundred μs . Rejection for events with time duration too large
 592 for an EAS signal, namely longer than 72 GTUs on the ISS, is also imple-
 593 mented. This implementation for one EC requires only a few per cent of the

resources of commercial FPGAs, which allows to implement it within the power constraints imposed on the ISS.

Tests performed with EUSO-Balloon and EUSO-TA data, as well as measurements performed at the TurLab facility, allowed validating the main functions of the algorithm. The system automatically adjusts the thresholds to keep the rate of triggers on background fluctuations below 1 Hz/EC even in the case of slow background variations. The FLT level trigger detects EAS-like events with light intensities comparable to those JEM-EUSO would observe in the energy range $E > 5 \cdot 10^{19}$ eV and in the presence of expected night sky background. These results strengthen those obtained in [7] and successive publications as they show that the trigger concept developed from simulation can be effectively implemented in hardware and performs well on real data.

The FLT has shown to be quite effective in rejecting city-like and other slow but bright events such as meteors. Of the few spurious triggers that occurred most were artefacts of discontinuities introduced by the available equipment.

The examples shown in this paper are only a sub-sample of all tests performed on the data reported here and the ongoing activities at the TurLab facility and EUSO-TA.

The VHDL logic of the FLT is currently being implemented on the FPGA of the PDM board. EUSO-SPB [26], the next stratospheric balloon flight, is expected to host this trigger logic on-board to verify its performance on real EAS.

6. Acknowledgements

This work was partially supported by the Basic Science Interdisciplinary Research Projects of RIKEN and JSPS KAKENHI Grant (22340063, 23340081, and 24244042), by the Italian Ministry of Foreign Affairs and International Cooperation, by the 'Helmholtz Alliance for Astroparticle Physics HAP' funded by the Initiative and Networking Fund of the Helmholtz Association, Germany, and by Slovak Academy of Sciences MVTs JEM-EUSO as well as VEGA grant agency project 2/0076/13. Russia is supported by the Russian Foundation for Basic Research Grant No 13-02-12175-ofi-m. The Spanish Consortium involved in the JEM-EUSO Space Mission is funded by MICINN & MINECO under the Space Program projects: AYA2009-06037-E/AYA, AYA-ESP2010-19082, AYA-ESP2011-29489-C03, AYA-ESP2012-39115-C03,

630 AYA-ESP2013-47816-C4, MINECO/FEDER-UNAH13-4E-2741, CSD2009-
631 00064 (Consolider MULTIDARK) and by Comunidad de Madrid (CAM) un-
632 der projects S2009/ESP-1496 & S2013/ICE-2822. The activities at TurLab
633 facility have been partially funded by the European High-Performance Infras-
634 tructures in Turbulence (EuHIT). The authors acknowledge strong support
635 from the French Space Agency CNES who provided - besides funding - the
636 leadership that made the EUSO-Balloon project possible in a very short time.
637 We are deeply indebted to the balloon division of CNES for a perfect launch,
638 smooth flight operation and flawless telemetry. The Canadian Space Agency
639 has provided outstanding facilities at the Timmins Stratospheric Balloon
640 Base, and a quick and careful recovery of the instrument. The authors dedi-
641 cate this paper to the memory of Dr. Jacek Karczmarczyk and Dr. Yoshiya
642 Kawasaki, who have contributed greatly to the project and will be deeply
643 missed.

644 References

- 645 [1] R. Benson and J. Linsley, Proc. 17th Int. Cosmic Ray Conf. (Paris),
646 (1981) 8.
- 647 [2] J.H. Adams et al. (JEM-EUSO Coll.), Exp. Astronomy 40 (2015) 3.
- 648 [3] F. Kajino et al. (JEM-EUSO Coll.), Nucl. Inst. Meth. A 623 (2010) 422.
- 649 [4] J.H. Adams et al. (JEM-EUSO Coll.), Exp. Astronomy 40 (2015) 19.
- 650 [5] M. Bertaina et al. (JEM-EUSO Coll.), Adv. Space Res. 53 (2014) 1515.
- 651 [6] C. Berat et al., Astrop. Phys. 33 (2010) 221.
- 652 [7] J.H. Adams Jr. et al. (JEM-EUSO Coll.), Astrop. Phys. 44 (2013) 76.
- 653 [8] M. Sato et al., Int. Journal of Mod. Phys. A 20/29 (2005) 6903.
- 654 [9] N. Sakaki et al., Proc. 30th Int. Cosmic Ray Conf. (Merida), HE 5 (2007)
655 1037.
- 656 [10] G.K. Garipov et al. (Tatiana Coll.), JETP Letters 4 (2005) 185.
- 657 [11] J. Bayer et al. (JEM-EUSO Coll.), Proc. 32th Int. Cosmic Ray Conf.
658 (Beijing), 3 (2011) 168; arXiv:1204.5065.

- 659 [12] J. Bayer et al. (JEM-EUSO Coll.), Proc. 33th Int. Cosmic Ray Conf.
660 (Rio de Janeiro), #0432 (2013); arXiv:1307.7071.
- 661 [13] M. Bertaina et al. (JEM-EUSO Coll.), Nucl. Instr. & Meth. A 824 (2016)
662 253.
- 663 [14] http://www.xilinx.com/support/documentation/data_sheets/ds150.pdf
- 664 [15] M. Bertaina et al. (JEM-EUSO Coll.), Exp. Astronomy 40 (2015) 117.
- 665 [16] J.H. Adams et al. (JEM-EUSO Coll.), Exp. Astronomy 40 (2015) 281.
- 666 [17] M. Bertaina et al. (JEM-EUSO Coll.), EPJ Web of Conferences 89
667 (2015) 03003.
- 668 [18] J.H. Adams et al. (JEM-EUSO Coll.), Exp. Astronomy 40 (2015) 301.
- 669 [19] <http://www.arduino.cc>
- 670 [20] P. von Ballmoos et al. (JEM-EUSO Coll.), Proc. 34th Int. Cosmic Ray
671 Conf. (Den Haag), #0725 (2015).
- 672 [21] M. Bertaina et al. (JEM-EUSO Coll.), Proc. 34th Int. Cosmic Ray Conf.
673 (Den Haag), #0890 (2015).
- 674 [22] J. Eser et al. (JEM-EUSO Coll.), Proc. 34th Int. Cosmic Ray Conf. (Den
675 Haag), #0860 (2015).
- 676 [23] F. Fenu et al. (JEM-EUSO Coll.), Proc. 34th Int. Cosmic Ray Conf.
677 (Den Haag), #0639 (2015).
- 678 [24] T. Abu-Zayyad, et al. (Telescope Array Coll.), Nucl. Instrum. Meth.
679 A689 (2012) 87.
- 680 [25] M. Casolino et al. (JEM-EUSO Coll.), Proc. 34th Int. Cosmic Ray Conf.
681 (Den Haag), #0636 (2015).
- 682 [26] L. Wiencke et al. (JEM-EUSO Coll.), Proc. 34th Int. Cosmic Ray Conf.
683 (Den Haag), #0165 (2015).